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Tsamis, Y. G.; Flores-Fajardo, N.; Henney, W. J.; Walsh, J. R. and Mesa-Delgado, A. (2013). Chemical abundances in Orion protoplanetary discs: integral field spectroscopy and photoevaporation models of HST 10. *Monthly Notices of the Royal Astronomical Society*, 430(4) pp. 3406–3421.

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Chemical abundances in Orion protoplanetary discs: integral field spectroscopy and photoevaporation models of HST 10^{*}

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Accepted 2013 January 21. Received 2013 January 18; in original form 2012 November 13

ABSTRACT

Photoevaporating protoplanetary discs (proplyds) in the vicinity of hot massive stars, such as those found in Orion, are important objects of study for the fields of star formation, early disc evolution, planetary formation and H II region astrophysics. Their element abundances are largely unknown, unlike those of the main-sequence stars or the host Orion nebula. We present a spectroscopic analysis of the Orion proplyd HST 10, based on integral field observations with the Very Large Telescope/FLAMES fibre array with 0.31×0.31 arcsec² spatial pixels. The proplyd and its vicinity are imaged in a variety of emission lines across a 6.8×4.3 arcsec² area. The reddening, electron density and temperature are mapped out from various line diagnostics. The abundances of helium, and eight heavy elements, are measured relative to hydrogen using the direct method based on the [O III] electron temperature. The abundance ratios of O/H and S/H are derived without resort to ionization correction factors. We construct dynamic photoevaporation models of HST 10 with the CLOUDY microphysics code that validate the oxygen and sulphur abundances. With the exception of [O I] $\lambda 6300$ and [S II] $\lambda 4069$, the model fit is satisfactory for all spectral lines arising from the proplyd. The models show that the classic ionization correction factor for neon significantly underestimates (0.4 dex) this element's abundance in the low ionization conditions of HST 10. Apart from iron, whose gas-phase abundance is ~ 0.3 dex lower than in the local Orion nebula, most other elements in the proplyd do not show substantially different gas-phase abundances from the nebula. The abundances of carbon, oxygen and neon in HST 10 are practically the same as those in B-type stars in Orion.

Key words: protoplanetary discs – stars: protostars – ISM: abundances – H II regions – ISM individual objects: HST 10 – ISM individual objects: Orion nebula.

1 INTRODUCTION

HST 10 (e.g. Johnstone, Hollenbach & Bally 1998) is an Orion proplyd of impressive appearance that lies farther away from the Trapezium than the compact group of bright ‘LV’ proplyds (Laques & Vidal 1979). It is thus subjected to a much smaller incident ionizing flux from the Trapezium stars. Fig. 1 shows a composite image of HST 10 based on the VLT data set presented in this paper

that highlights its ‘teardrop’ morphology and the presence of both neutral and ionized gas emission. The proplyd is designated as 182–413 according to the nomenclature of Orion sources by O’Dell & Wenn (1994). It is larger and fainter than the proplyds found close to the Trapezium, with a less elongated and less symmetric tail. This is in line with the general trends seen in these sources (Bally et al. 1998; O’Dell 1998), which can be understood in terms of a model whereby protostellar discs around the young low-mass stars in the nebula are evaporated by the ultraviolet radiation from the high-mass Trapezium stars (Johnstone et al. 1998; Henney & Arthur 1998; Störzer & Hollenbach 1999; Vasconcelos et al. 2005).

In this study, we present optical integral field spectroscopy (IFS) of HST 10 and its immediate Orion environment, constructing

^{*} Based on observations made with ESO telescopes at the Paranal Observatory under programme 078.C-0247(A).

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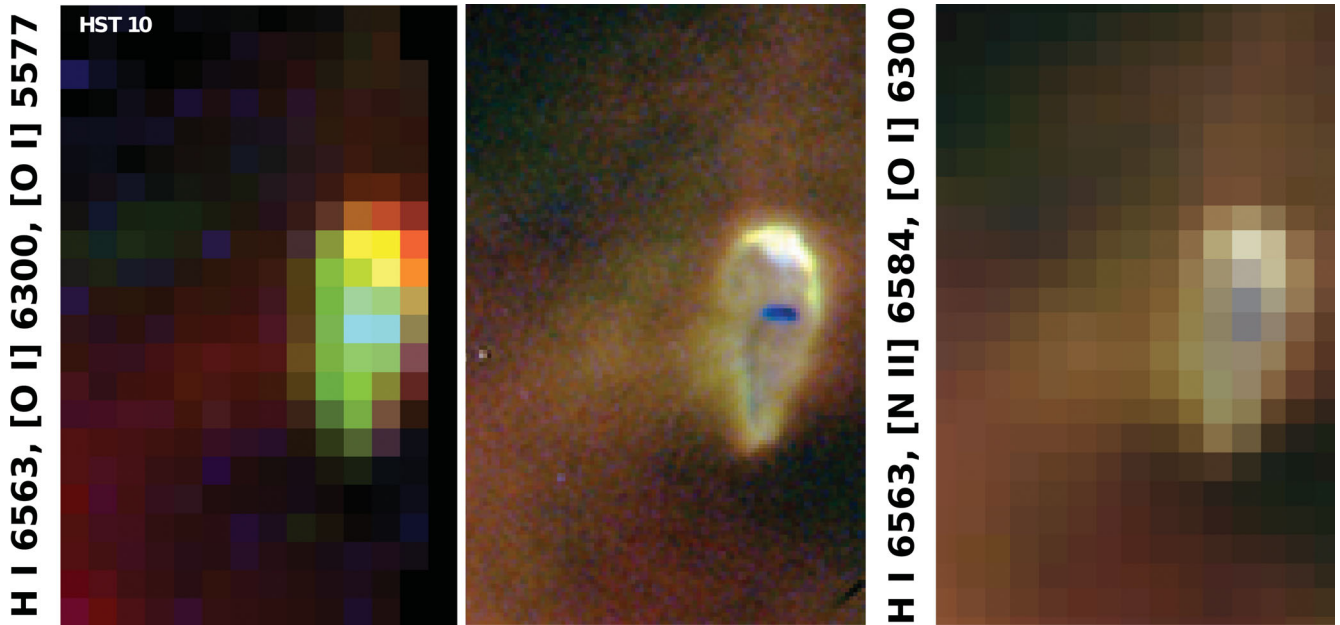


Figure 1. Comparison of VLT versus *HST* images of the HST 10 proplyd: (left) composite of VLT Argus monochromatic channel maps with $H\alpha$ $\lambda 6563$ (red), $[O\text{I}]$ $\lambda 6300$ (green), $[O\text{I}]$ $\lambda 5577$ (blue). The embedded circumstellar disc dominates the emission of $[O\text{I}]$ $\lambda 5577$. The field of view is 6.8×4.3 arcsec²; (centre): *HST* composite using filters F656N (red), F658N (green), F631N (red); (right) the *HST* image convolved with a Gaussian filter of 0.8 arcsec full width at half maximum (approximately equal to the seeing during the VLT observations) and rebinned to the VLT pixel scale. North is up and east is to the left-hand side.

photoevaporation models of the proplyd. Our aim is to enlarge the number of sources with element abundance determinations following the studies of LV 2 (Tsamis et al. 2011; Tsamis & Walsh 2011) and HST 1 (Mesa-Delgado et al. 2012) proplyds.

1.1 HST 10 in the Orion nebula context

The location on the sky of HST 10 is halfway between the inner Trapezium cluster and the Bright Bar region (see Fig. 2), at an angular separation of about 1 arcmin to the south-east from θ^1 Ori C (O7V spectral type; e.g. Simón-Díaz et al. 2006), the dominant illuminating star of the Orion nebula. Kinematic studies of the emission from the proplyd (Henney & O’Dell 1999) suggest that it is situated in the foreground of the nebula, with a true separation from θ^1 Ori C of 0.2–0.3 pc. Unlike in many of the smaller, brighter proplyds, the circumstellar disc of mass $\sim 0.0054 M_{\odot}$ (see Section 4.3) is clearly visible in HST 10. In most emission lines, the disc is seen in absorption, but it is seen in emission in the H_2 $2.12 \mu\text{m}$ line (Chen et al. 1998), the $[O\text{I}]$ $\lambda 6300$ line (Johnstone et al. 1998) and the $[O\text{I}]$ $\lambda 5577$ line (our Fig. 1). The $\lambda 6300$ line emission is shown in red in the right-hand panel of Fig. 2. There is also a faint high-velocity microjet detected in $[O\text{I}]$ $\lambda 6300$ (Henney & O’Dell 1999), which extends perpendicular to the disc (Bally, O’Dell & McCaughrean 2000).

Although the proplyds close to the Trapezium tend to be very symmetric about the line that joins them to θ^1 Ori C, this is not so true of HST 10, which seems to be governed by two different axes. The rotational axis of the disc and jet is oriented north–south, while the direction to θ^1 Ori C is at roughly 30° to this, at a position angle of 330° . This produces considerable distortion in the shape of the proplyd, with the tail, in particular, seemingly more governed by the jet axis than the ionizing radiation field.

The Orion nebula is highly structured at all scales and many previously studied nebular features pass through or near the immediate

vicinity of HST 10, as illustrated in Fig. 2. The shadow rays (O’Dell 2000) are low-ionization linear features, which are seen outside the positions of some of the brighter proplyds, and are exactly aligned with the outer segment of the line joining θ^1 Ori C and the proplyd (see fig. 2 of O’Dell et al. 2009). Three such rays are indicated in Fig. 2: the two most prominent, which are cast by 177–341 (HST 1) and 159–350 (HST 3), plus a much fainter one, which is cast by 170–337 and which passes within 2 arcsec of HST 10 in projection. Unlike HST 10, 170–337 shows evidence of being located behind the Trapezium stars (Henney & O’Dell 1999), and therefore its shadow must be also; therefore it is very unlikely to be physically close to HST 10.

The bar features are another type of linear structure that is very common in the nebula of which the Bright Bar is the most prominent example (e.g. O’Dell & Yusef-Zadeh 2000; Mesa-Delgado et al. 2012; Rubin et al. 2011). These are regions where the line of sight is tangential to a local ionization front, but the exact geometry is unclear in many cases. García-Díaz & Henney (2007) found that some faint compact bars are associated with dark lanes that are seen as linear extinction features in redshifted velocity channels. Such is the case for the jagged low ionization filament that crosses the field of HST 10, as shown in the right-hand panel of Fig. 2. A section of this filament has been detected by VLT Argus and its electron temperature and density have been determined from our spectra. This type of emission feature may represent the ionized skin of a dense molecular filament that is protruding into the $H\text{II}$ region, in which case it is likely to be close to the principal ionization front in the background of the nebula. Therefore again, the line-of-sight position of the feature is probably far from HST 10.

A different type of linear structure is the high velocity collimated jets that drive the numerous Herbig–Haro bow shocks seen in the nebula. The driving jet of the HH 203 bow shock (Doi, O’Dell & Hartigan 2004) passes within 5 arcsec of HST 10, although it is just outside the field of view of our VLT observations. A much larger

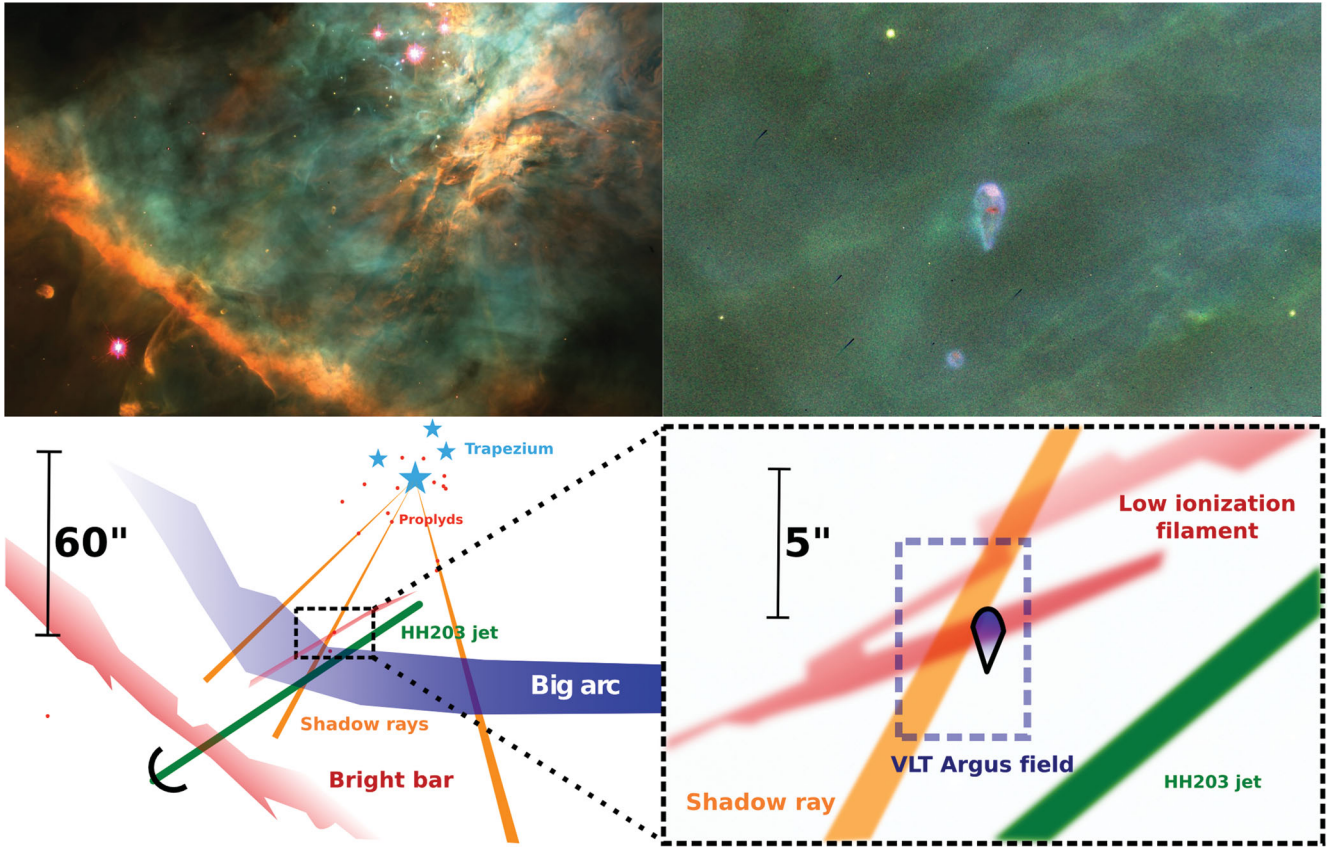


Figure 2. HST 10 in the context of the Orion nebula (to be viewed anti-clockwise). Left upper panel: large-scale, three-colour image of the Orion nebula in the filters of [N II] (red), H α (green) and [O III] (blue), based on a mosaic of *HST* WFC observations described in O’Dell & Wong (1996). Left lower panel: schematic view of various components in Orion. Right lower panel: sketch of the HST 10 field (dashed cyan box) and its immediate environs. Right upper panel: three-colour image centred on HST 10 in the filters of [O I] (red), [S II] (green) and [N II] (blue), based on *HST* PC observations described in O’Dell (1998). The labelled objects are described in the text.

scale kinematic feature is the so-called Big Arc (Doi et al. 2004; García-Díaz & Henney 2007), which is a blueshifted high-ionization structure that extends over several arcmin and whose origin is unclear. Although HST 10 is very close to the northern boundary of this feature, it probably does not affect our observations.

2 THE VLT OBSERVATIONS AND DATA REDUCTION

IFS of HST 10 ($\alpha_{J2000} = 5^h35^m18^s.204$, $\delta_{J2000} = -5^\circ24'13''.38$) was performed on the 8.2-m VLT/UT2 Kueyen at the European Southern Observatory (ESO) on Paranal during two nights in 2006 October and December with the fibre-fed FLAMES Giraffe spectrograph (Pasquini et al. 2002) and the Argus microlens array in high magnification mode. A field of view of 6.8×4.3 arcsec² was observed and 297 positional spectra were recorded in the optical range (3620–7184 Å) using a 2K \times 4K EEV CCD (with pixels of 15 μ m size). Six low-resolution (LR) grating settings were employed (Table 1). The size of the angular resolution element was 0.31×0.31 arcsec², corresponding to a spatial scale of 127×127 au² at the distance to M42 (410 pc; Reid et al. 2009). The measured seeing during the observations with the various gratings was 0.7 (LR1), 0.7–0.9 (LR2), 0.9 (LR3), 0.8–0.9 (LR4), 0.7–0.8 (LR5) and 0.7 arcsec full width at half maximum (LR6).

Due to a slight misalignment of the guide/fiducial star fibres used for the acquisition of the Argus field of view, the proplyd is not

Table 1. Journal of VLT/FLAMES Argus observations.

Date (UT)	λ -range (Å)	Grating	$\lambda/\delta\lambda$	Exp. time (s)
2006/12/15	3620–4081	LR1	12 800	2×150
2006/12/15	3964–4567	LR2	10 200	3×390
2006/12/15	4501–5078	LR3	12 000	4×206
2006/10/03	5015–5831	LR4	9 600	3×254
2006/10/03	5741–6524	LR5	11 800	3×255
2006/10/03	6438–7184	LR6	13 700	4×179

exactly centred in the array and is shifted towards the western edge of the field (Fig. 1, left-hand panel). In Fig. 1 (right-hand panel) we show a *Hubble Space Telescope* (*HST*) image of the proplyd that has been degraded to the resolution of the Argus spaxels, taking into account the seeing measured during our VLT observations. Comparison of these two accurately aligned images shows that no appreciable loss of flux has resulted from the proplyd based on its position within the Argus field of view. The resolution achieved allows us to clearly distinguish three regions in HST 10: cusp, core/disc area and tail. We cannot, however, very well resolve the structure within these.

The 2D data arrays were cosmic ray cleaned, flat-fielded, wavelength calibrated and extracted utilizing the girBLDRS pipeline developed by the Geneva Observatory (Blecha & Simond 2004).

The flux calibration was done within IRAF using contemporaneous exposures of various spectrophotometric standards for the grating settings: EG 21 (LR1), Feige 110 (LR2, LR3), LTT 7987 (LR4, 5, 6). The reduction products consisting of separate signal and error arrays were converted into (X, Y, λ) data cubes using IRAF routines and the Argus fibre positioning table.

The data cubes were corrected for the effects of differential atmospheric refraction (DAR) using the algorithm of Walsh & Roy (1990) as in Tsamis et al. (2011). In this way, the various emission line or continuum maps were registered on an identical spatial grid across the entire spectral range of Table 1, allowing us to form high fidelity ratio maps of spectral features widely separated in wavelength. The shifts involved in the correction for DAR were extremely small as the airmasses during the observations were low (<1.3). There was a detectable shift in the positioning of Argus between the two nights and as a result the LR4, 5, 6 spatial grid was shifted by one spaxel due east to match the LR1, 2, 3 cubes. The necessary rebinning has contributed to the edge effects apparent on the left-hand side of some of the maps discussed in Section 3.1. The LR5 and LR6 spectra were scaled by a factor of 0.813 via the flux of the [N II] $\lambda 5755$ line which falls on both LR4 and LR5 gratings in order to match the bluer spectral coverage. The LR1 spectrum was scaled to the LR2 spectrum using the H I $\lambda 3970$ line common to both gratings.

We selected emission lines of interest in the various grating setups and fitted them with Gaussian profiles in all ~ 300 individual positional spectra. An automated procedure was used based on MIDAS tasks utilizing the Newton–Raphson least-squares method. The local continuum bracketing each line was fitted using a first order polynomial or a cubic spline. The resulting arrays were further processed within MIDAS and IRAF to form maps of the lines or continua and their signal-to-noise ratios.

3 MAPPING THE PHYSICAL DIAGNOSTICS

3.1 Emission-line maps

In Figs 3 and 4, a collection of spectral maps of the HST 10 field is presented. These represent the observed line fluxes measured via Gaussian fitting as discussed above. The good seeing has resulted in a very good image quality, especially during the observations with the LR4, 5, 6 gratings. The proplyd shows an elongated morphology along the north–south direction resembling closely the higher resolution imagery of the *HST*. A filamentary feature extending towards the southeast is evident in several lines, but is mostly prominent in H I, He I, [N II] $\lambda 6584$, and particularly in [S II] $\lambda \lambda 6716, 6731$ emission. This is probably the ionized skin of a background molecular filament that is protruding into the H II region as discussed in Section 1.1.

All lines, with the prominent exception of [O I] $\lambda 5577$ and [N I] $\lambda \lambda 5198, 5200$,¹ peak on the upper half (the ‘cusp’) of the proplyd. The lines from neutral species, particularly [O I] $\lambda 5577$ (Fig. 1), peak at spaxels coincident with the position of the embedded circumstellar disc which has been detected in narrow band H α absorption, and [O I] $\lambda 6300$ (Bally et al. 1998) and molecular hydrogen emission (Chen et al. 1998). These are the first detections of HST 10 in several other tracers, including the auroral [O III] $\lambda 4363$, [S II] $\lambda 4069$ and [N II] $\lambda 5755$ lines; [Fe III] $\lambda 5270$; [Ar III]

$\lambda 7136$; and of the disc in [O I] $\lambda 5577$. The C II $\lambda 4267$ pure recombination line (RL) (Davey, Storey & Kisielius 2000) is also detected from the cusp and is used in Section 4 as a C²⁺ abundance diagnostic.

The H α and [O III] $\lambda 4959$ maps show local minima at the position of the [O I] $\lambda 5577$ peak emission, consistent with the underlying disc being shielded from Lyman continuum photons by the ionization front of the outer cusp. The [S II] red doublet lines show peaks on both the cusp and the tail of HST 10. Most of the forbidden lines detected here are collisional lines (CLs) produced when low-lying ionic levels are thermally excited by electron impacts. Exceptions are the [O I] lines which, near the disc, are emitted when the photo-dissociation of OH at the base of a neutral evaporated flow arising from the disc produces electronically excited atomic oxygen (Störzer & Hollenbach 1998). On the other hand, at the proplyd’s cusp, the emission of [O I] $\lambda 6300$ is probably dominated by electronic collisional excitation.² Also, the [N I] doublet lines are emitted in atomic gas absorbing far-UV stellar radiation in a fluorescence process (Ferland et al. 2012). A similar fluorescence process may be affecting the C II $\lambda 6578$ line (Fig. 4) whose utility as an abundance diagnostic is limited (e.g. Esteban et al. 2004).

3.2 Reddening and ionization degree

The interstellar reddening in the line of sight towards HST 10 was determined from a comparison of the observed H α /H β , H δ /H β , H I $\lambda 3970$ /H β ratios with their predicted Case B recombination theory ratios from Storey & Hummer (1995). We used the modification by Blagrove et al. (2007) of the reddening law of Cardelli, Clayton & Mathis (1989) with a total to selective extinction ratio $R_V = 5.5$ applicable to M42. A map of the resulting logarithmic $c(H\beta)$ reddening parameter is shown in Fig. 5, where the individual results from the above line ratios were given weights of 3:1:1. The individual line ratios over the proplyd’s cusp result in these mutually consistent $c(H\beta)$ values: 0.52 ± 0.10 (H α /H β), 0.49 ± 0.15 (H δ /H β) and 0.60 ± 0.11 (H I $\lambda 3970$ /H β); these are mean values over the nine spaxels of our designated cusp area. Overall, the reddening appears to be higher over HST 10 and shows values of ≈ 0.6 at the position of the peak H α emission. Our reddening determination is in good agreement with the value obtained from *HST* image analysis by O’Dell (1998), and by O’Dell & Yusef-Zadeh (2000) who calculated $c(H\beta)$ from the 20 cm/H α ratio. The latter study shows that there is a general trend of increasing reddening (0.45–0.60) from south to north across an area coincident with the Argus field of view, while our reddening map of Fig. 5 shows ~ 0.4 –0.55. The spatial variation is almost identical between the two techniques, whereas the absolute values of the extinction at H β derived by O’Dell & Yusef-Zadeh are 12 per cent higher. Adopting such a higher extinction value would reduce by 3 per cent the ratios with respect to H β of lines in the red portion of the spectrum, which is smaller than the observational uncertainty in our measurements (Table 2).

The enhanced $c(H\beta)$ at the proplyd position is not likely due to grains mixed in with the ionized gas. This is because the dust optical depth through the ionized gas is proportional to the ionization parameter, which for the proplyd is quite low. Our model in

¹ The mean flux from both [N I] lines is shown in Fig. 3.

² The [O I] line maps shown here contain a night-sky component which cannot be subtracted as the Argus sky allocated fibres could not be pointed to regions free from Orion nebula emission.

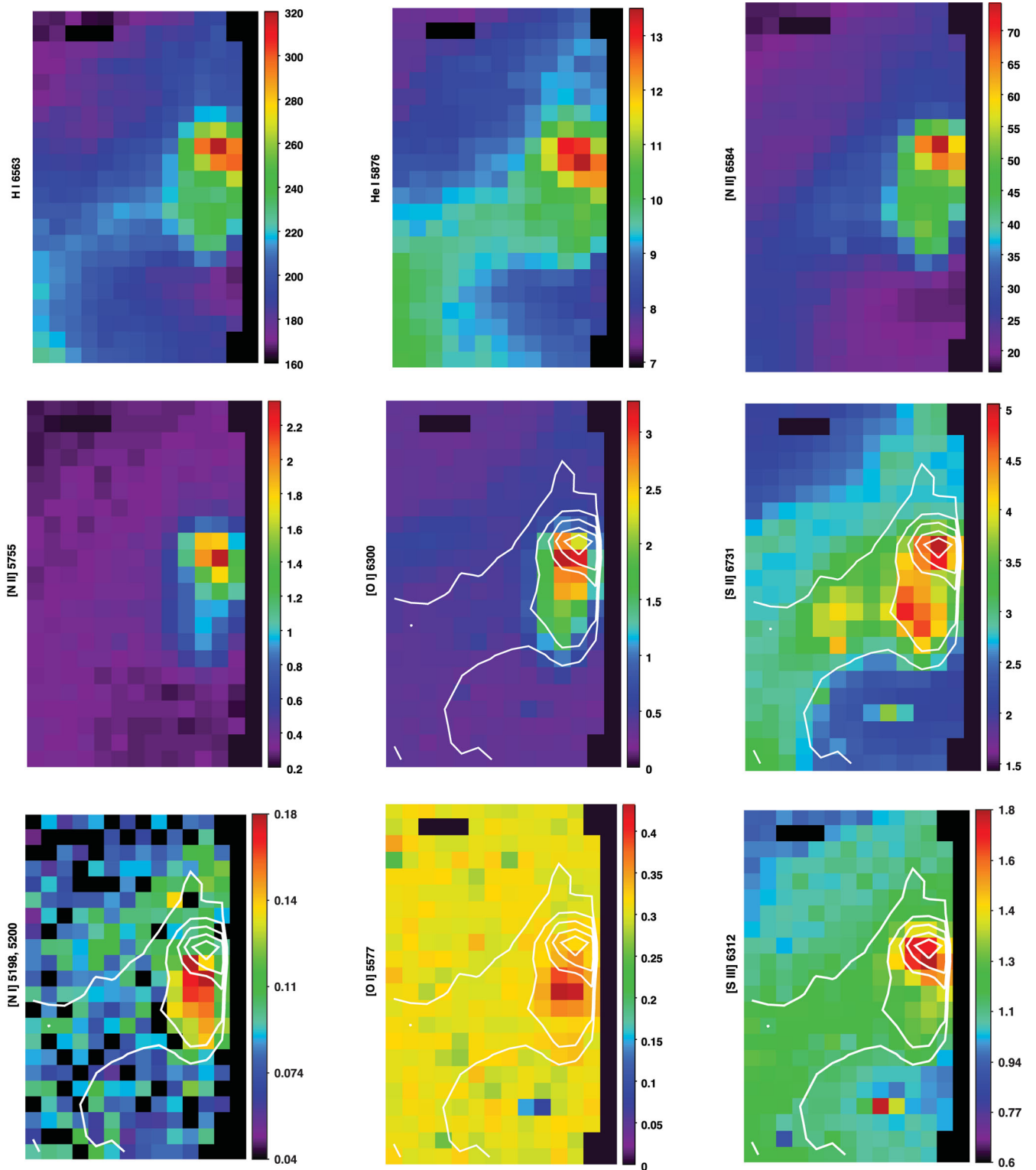


Figure 3. Emission line maps (in units of $10^{-15} \text{ erg s}^{-1} \text{ cm}^{-2} \text{ spaxel}^{-1}$, not corrected for reddening; the isocontours are the $H\alpha$ line). The field of view is $6.6 \times 4.2 \text{ arcsec}^2$ (22 by 14 spaxels). Dark masked outer columns are artefacts at the edges of the array introduced by the rebinning used to correct for DAR and for the spatial shift between the LR1, 2, 3 and LR4, 5, 6 observations (see the text for details). Two spaxels in the fourth row are affected by a CCD defect and three blank spaxels in row 21 correspond to dead fibres. North is up and east is to the left-hand side.

Section 4.2 indicates a V-band optical depth of only about 0.05 to the ionization front (see Fig. 10). Therefore, most of the extra extinction compared to the Orion nebula must be due to dust between the ionization front and the surface of the embedded circumstellar

disc, which should have A_V of the order of unity. This is further supported by geometric arguments as the proplyd's axis of symmetry is tilted with respect to the line of sight by 150° with the cusp facing away from the observer (Table 5). As a result, emission from the

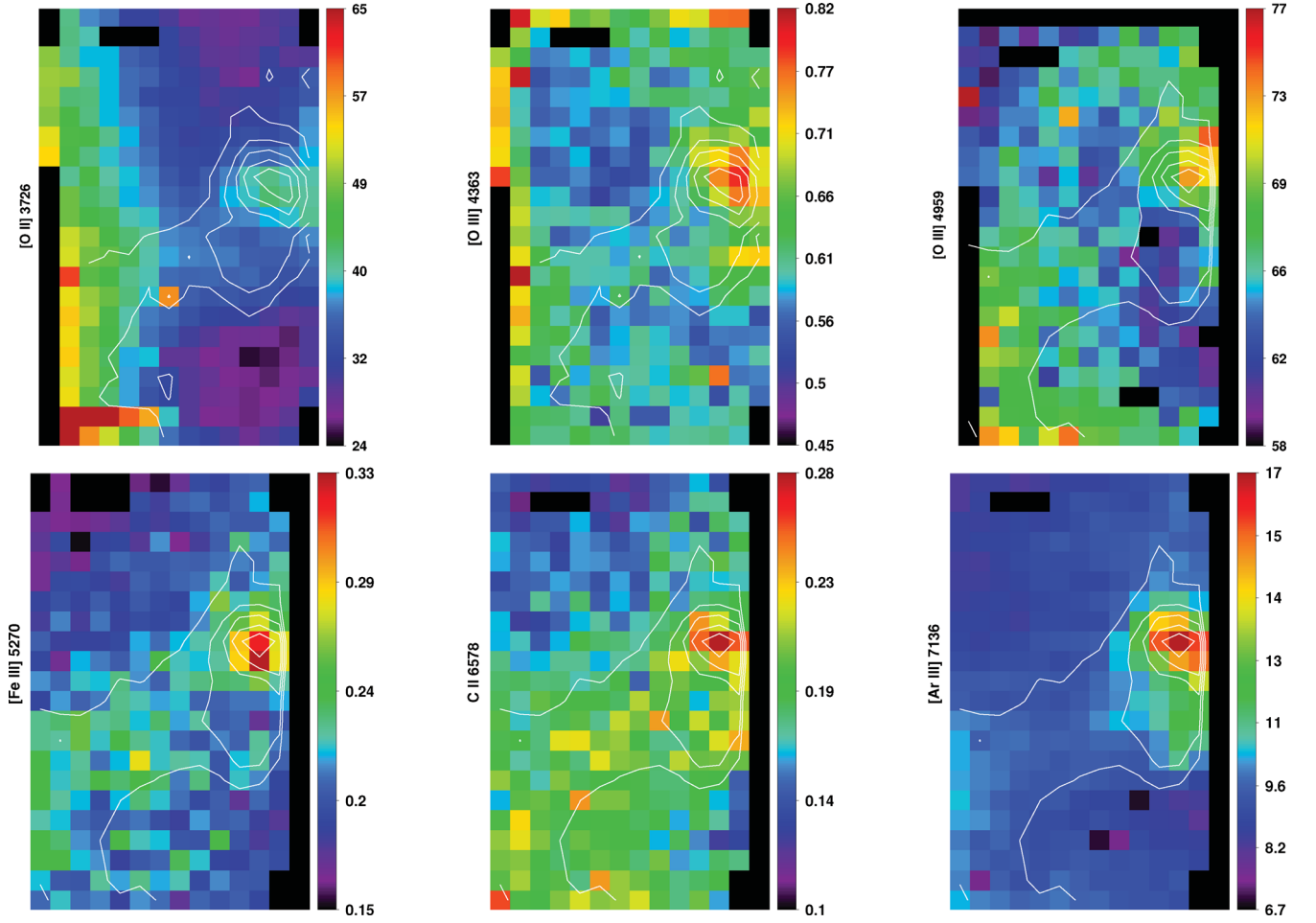


Figure 4. As in the previous figure. The contours are H α , except on the [O II] λ 3726 and [O III] λ 4363 maps which show contours from H δ .

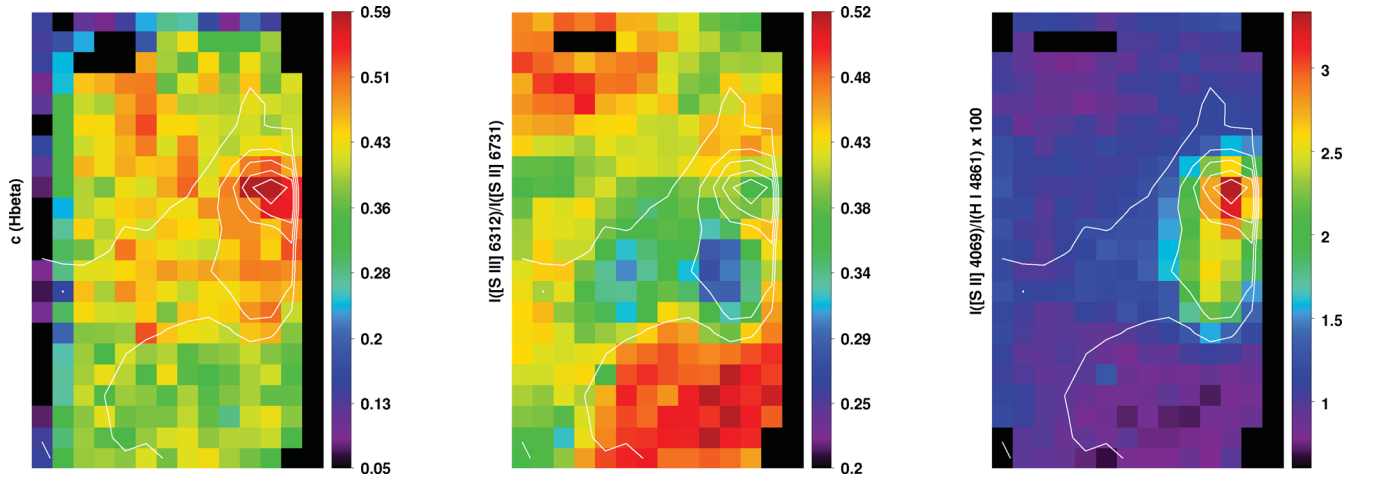


Figure 5. The reddening parameter $c(\text{H}\beta)$ and a sample of dereddened maps: the [S III] λ 6312/[S II] λ 6731 ratio and the [S II] λ 4069/H β ratio.

cusps suffers reddening as it passes through the intervening outflow arising from the disc.

The lower part of the Argus field containing Orion nebula emission shows lower $c(\text{H}\beta)$ values, by ~ 0.3 dex. Such moderate reddening is in overall agreement with other studies of the Orion H II region which focused on Orion nebula emission (O'Dell & Yusef-Zadeh 2000; Rubin et al. 2003; Esteban

et al. 2004; Blagrove et al. 2007; Tsamis et al. 2011), proplyds (O'Dell 1998; Tsamis et al. 2011; Mesa-Delgado et al. 2012), and Herbig–Haro objects (Mesa-Delgado et al. 2009; Núñez-Díaz et al. 2012).

The above $c(\text{H}\beta)$ map was used to convert the observed line maps to dereddened intensity maps relative to $I(\text{H}\beta)$, such that $I(\lambda)/I(\text{H}\beta) = F(\lambda)/F(\text{H}\beta) \times 10^{c(\text{H}\beta)f(\lambda)}$, where $f(\lambda)$ is the Blagrove

Table 2. Dereddened line fluxes from the (background-subtracted) cusp of HST 10 and the local M42 background (in units such that $H\beta = 100$).

Line	HST 10 $I(\lambda)$	M42 $I(\lambda)$
$c(H\beta)$	0.53 ± 0.04	0.42 ± 0.06
[O II] $\lambda 3726$	89.0 ± 17.9	73.3 ± 8.0
[Ne III] $\lambda 3967$	2.54 ± 0.30	5.21 ± 0.27
H I $\lambda 3970$	13.3 ± 0.38	15.8 ± 0.67
[S II] $\lambda 4069$	4.45 ± 0.47	0.875 ± 0.056
H δ $\lambda 4101$	22.1 ± 0.55	26.5 ± 0.72
C II $\lambda 4267$	0.122 ± 0.048	0.275 ± 0.027
H γ $\lambda 4340$	40.4 ± 4.6	49.6 ± 0.67
[O III] $\lambda 4363$	0.796 ± 0.181	1.16 ± 0.03
He I $\lambda 4471$	3.81 ± 0.72	5.05 ± 0.09
O II $\lambda 4649$	–	0.130 ± 0.039
[Ar IV] $\lambda 4711$	–	0.059 ± 0.030
[Ar IV] $\lambda 4740$	–	0.066 ± 0.033
$H\beta$ $\lambda 4861$	100.0	100.0
[O III] $\lambda 4959$	55.2 ± 9.4	117.5 ± 3.6
[O I] $\lambda 5577$	0.193 ± 0.033	0.495 ± 0.116
[Fe III] $\lambda 5270$	0.297 ± 0.039	0.330 ± 0.037
[Cl III] $\lambda 5518$	0.172 ± 0.039	0.400 ± 0.083
[Cl III] $\lambda 5538$	0.267 ± 0.038	0.480 ± 0.085
[N II] $\lambda 5755$	2.32 ± 1.16	0.470 ± 0.049
He I $\lambda 5876$	10.3 ± 2.1	13.9 ± 0.3
[O I] $\lambda 6300$	2.82 ± 2.08	0.535 ± 0.128
[S III] $\lambda 6312$	1.36 ± 0.39	1.74 ± 0.46
H α $\lambda 6563$	235.6 ± 2.2	276.5 ± 7.9
[N II] $\lambda 6584$	72.1 ± 6.2	31.7 ± 1.4
C II $\lambda 6578$	0.176 ± 0.033	0.350 ± 0.024
[S II] $\lambda 6716$	1.82 ± 0.32	2.44 ± 0.60
[S II] $\lambda 6731$	4.00 ± 0.67	3.63 ± 0.88
[Ar III] $\lambda 7136$	13.1 ± 1.38	12.1 ± 1.5
$H\beta$ $\lambda 4861^a$	148.7 ± 51.1	142.0 ± 21.0
$T_e([O III])$ (K)	9070 ± 240	8390 ± 100
$N_e([S II])$ (cm^{-3})	21560 ± 6227	2860 ± 400
$N_e([N II])$ (cm^{-3})	46785 ± 32281	–

^a In units of $10^{-15} \text{ erg s}^{-1} \text{ cm}^{-2} \text{ spaxel}^{-1}$. The cusp area comprises nine spaxels and the background area ~ 20 spaxels.

et al. (2007) reddening curve. The dereddened intensity ratio maps were used for the subsequent physical analysis of the HST 10 field.

In Fig. 5 also the $[S III]/[S II]$ ratio is shown, which is a good indicator of the variation of the ionization degree across the field. HST 10 is a relatively low ionization structure overall. The ionization degree peaks in spaxels free from HST 10 emission in the southern part of the field. On HST 10, it shows distinct minima in the tail of the proplyd, southeast of the embedded disc, as well as across an extended area directly east of the proplyd. The former low ionization area roughly coincides with an area of polycyclic aromatic hydrocarbon emission imaged with adaptive optics at the Keck telescope (Kassis et al. 2007). The latter feature, which also displays relatively higher reddening than the nebula emission to the southwest, is the background $[S II]$ - and $[N II]$ -bright filament (Section 1.1).

3.3 Electron density and temperature

There is a variety of plasma diagnostics in the current data set for the determination of the electron density, N_e , and temperature, T_e . The TEMDEN task in IRAF v2.14 was employed. Maps of N_e are

shown in Fig. 6, based on the $[S II]$ $\lambda 6731/\lambda 6716$ doublet ratio and the $\lambda 4069/\lambda 6731$ ratio. In both cases, the overall morphology of the proplyd is well reproduced by the observed density distribution. Based on the former ratio, the mean $N_e([S II])$ is $4170 \pm 730 \text{ cm}^{-3}$ over an area within the 80 per cent $H\alpha$ isophote. The $\lambda 6731/\lambda 6716$ ratio, however, approaches the high density limit dictated by the low critical densities ($1200, 3300 \text{ cm}^{-3}$) of the $[S II]$ doublet upper levels, and its measured variation in the cusp of the proplyd indicates densities from $\sim 3000 \text{ cm}^{-3}$ to a poorly defined upper value. The $\lambda 6731/\lambda 6716$ density map is therefore only shown for illustration as it is *not accurate* over the proplyd area. We thus also employ the $[S II]$ $\lambda 4069/\lambda 6731$ ratio which is sensitive to densities of up to $\sim 2 \times 10^6 \text{ cm}^{-3}$. The corresponding map, which has been computed from the background-subtracted ratio over an area of 32 spaxels covering HST 10, shows peak densities of $3.1 \times 10^4 \text{ cm}^{-3}$ in the cusp of the proplyd.

For comparison, the proplyd density can also be estimated from the $H\alpha$ surface brightness as follows. From the peak $H\alpha$ emission, the dereddened brightness is $S(H\alpha) = 1.475 \times 10^{11} \text{ photon s}^{-1} \text{ cm}^{-2} \text{ sr}^{-1}$, whereas the background shows $8.53 \times 10^{10} \text{ photon s}^{-1} \text{ cm}^{-2} \text{ sr}^{-1}$. The emission measure (EM) is $EM = N_e^2 dL$ (where dL is the photon path length), or $4\pi S(H\alpha)/\alpha_{\text{eff}}(H\alpha)$. Substituting for the background-subtracted brightness and assuming a cusp radius or path length equal to two spaxels ($= 2 \times 0.31 \times 440 \text{ au}$), with $\alpha_{\text{eff}}(H\alpha) = 10^{-13} \text{ cm}^3 \text{ s}^{-1}$ at $10\,000 \text{ K}$, then $N_e = 4.45 \times 10^4 \text{ cm}^{-3}$. This is in fair agreement with the direct $[S II]$ $\lambda 4069/\lambda 6731$ ratio measurement.

There is an area of increased density directly east of the proplyd that corresponds to a lower ionization area visible on the $[S III]/[S II]$ ratio image. The N_e of this filamentary feature is $3200 \pm 300 \text{ cm}^{-3}$ over 20 spaxels based on the $\lambda 6731/\lambda 6716$ ratio and 3030 ± 340 based on the $\lambda 4069/\lambda 6731$ ratio. The agreement between the N_e determinations from the two ratios over spaxels free of HST 10 emission lends support to the validity of both diagnostics in the lower density regime of the local nebula. The peak density over this filament is up to a factor of ~ 2 higher than over the background nebula in the southwest corner of the field. This is most likely a background filament projected near HST 10's position as mentioned in Section 1.1.

The T_e distribution is shown in the last panel of Fig. 6 based on the background-subtracted $[O III]$ $\lambda 4363/\lambda 4959$ ratio over the HST 10 area. The temperature over the proplyd is $8740 \pm 290 \text{ K}$ within the 80 per cent $H\alpha$ isophote (28 spaxels). The $[S II]$ $\lambda 4069/\lambda 6731$ density map has been adopted for the computation. The cusp of HST 10 shows $T_e = 9070 \pm 240 \text{ K}$, while the tail has $8620 \pm 120 \text{ K}$. The proplyd cusp is of higher temperature than the local M42 background that has $T_e = 8390 \pm 100 \text{ K}$. In Fig. 7 we show a diagnostic diagram of the N_e - T_e plane applicable to the cusp of HST 10. It is based on background-subtracted line ratios and demonstrates that the $[S II]$ $\lambda 4069/\lambda 6731$ ratio is a fairly accurate density diagnostic, displaying a variation of 28 per cent about the mean at $T_e = 9070 \pm 240$. On the other hand, the $[N II]$ ratio is substantially more sensitive to the temperature as well, showing a large variation of 69 per cent about the mean in the same T_e range (Table 2), and cannot be relied upon. In the abundance analysis that follows the $[S II]$ $\lambda 4069/\lambda 6731$ densities and $[O III]$ $\lambda 4363/\lambda 4959$ temperatures will be therefore adopted.

In Table 2 the mean line fluxes and physical conditions over the bright cusp of HST 10 as well as over spaxels free from proplyd emission are listed. In Column 2 of Table 2 the values shown correspond to background-subtracted quantities following a straightforward extrapolation procedure, according to which spaxels close to

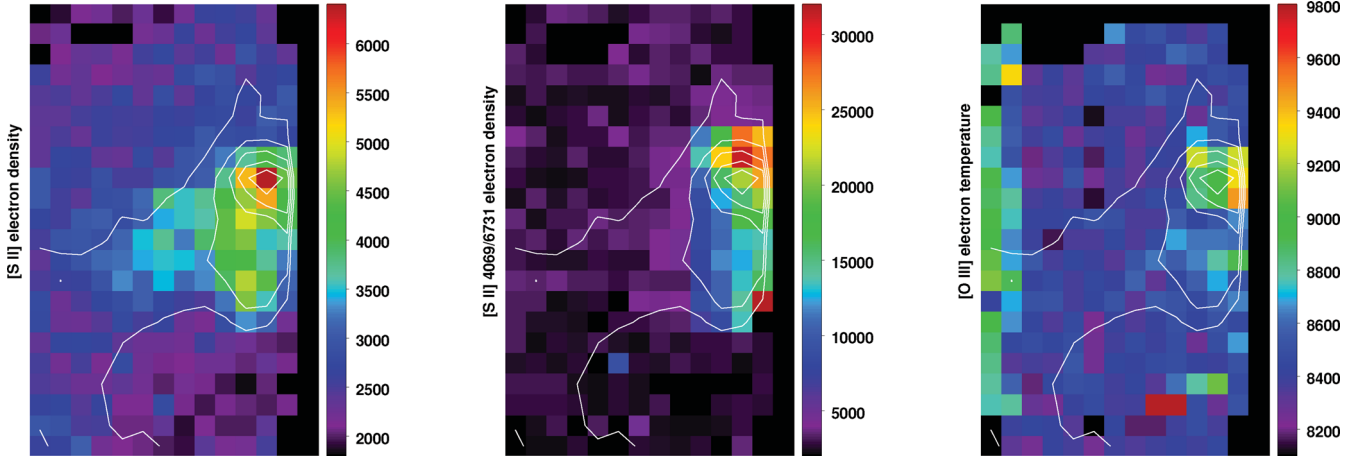


Figure 6. Electron density from the [S II] $\lambda 6716/\lambda 6731$ and $\lambda 4069/\lambda 6731$ ratios (in cm^{-3}), and electron temperature from the [O III] $\lambda 4363/\lambda 4959$ ratio (in kelvin). The $\lambda 4069/\lambda 6731$ and $\lambda 4363/\lambda 4959$ diagnostic maps are based on background-subtracted fluxes over spaxels showing HST 10 emission.

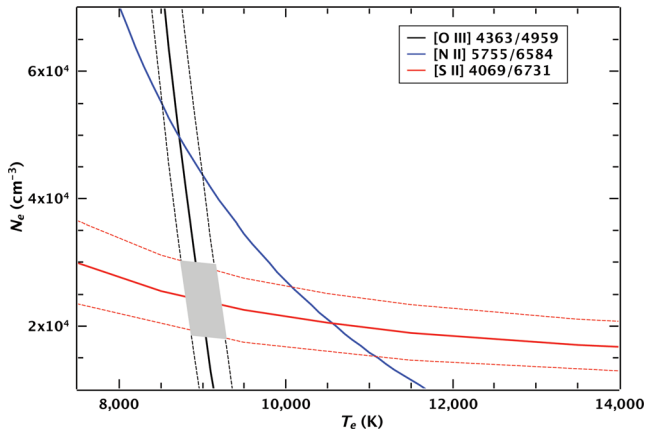


Figure 7. Diagnostic diagram applicable to mean line ratio values (solid lines) over the cusp of HST 10 (nine spaxels): $(X, Y) = (13-15, 2-4)$ in Argus array coordinates, where X, Y are, respectively, the long and short axes with a point of origin at the bottom right corner of the maps in Fig. 3. Dashed lines bracketing solid lines of the same colour denote the 1σ range of the corresponding ratios. The equivalent range for the [N II] diagnostic is too large (see Table 2) and is not shown (the same holds for the [S II] $\lambda 6731/\lambda 6716$ ratio which is not included in this plot; see the text for details). The shaded area corresponds to the adopted conditions for the empirical analysis of the cusp in Section 4.1.

the proplyd have been chosen as representative of the mean nebular background in front of and behind the source. The Orion nebula background shows appreciable structure down to scales of just a few spaxels (<1 arcsec). The scarcity of suitable background spaxels to the west of the proplyd might have somewhat affected the accuracy of our method. Inspection of *HST* images (Fig. 2, right-hand panels) shows that to the west of HST 10, immediately outside the Argus field of view, the background appears as a mix of featureless nebulosity and the low ionization filament which passes behind the proplyd as discussed in Section 1.1. Features with similar characteristics are present within the Argus field covered by spaxels below and to the left/upper left of the proplyd and they have been included in our characterization of the background.³ Therefore, the

location and number of spaxels used in the determination of the background in the proplyd's close vicinity are satisfactory for our purposes.

The quoted uncertainties in Table 2 are associated with the measured variation of a given quantity across the sampled spaxels (the rms deviation from the mean). The formal uncertainties, propagated using the reduction pipeline-generated error arrays, on the density or temperature-sensitive line ratios that are important for the abundance determinations are much smaller (less than two per cent).

4 CHEMICAL ELEMENT ABUNDANCES

We employ two complementary approaches for determining the gas-phase elemental abundances. The first is the traditional semi-empirical method (e.g. Seaton 1960), which in its simplest form assumes that each ion is emitted from a homogeneous zone with a constant electron density and temperature, and employs ionization correction factors (ICFs; e.g. Peimbert & Costero 1969; Kingsburgh & Barlow 1994) to account for unobserved ionization stages. The second is to construct self-consistent physical models of the proplyd ionization front and photoevaporation flow (e.g. Henney & O'Dell 1999; Henney et al. 2002; Henney, Arthur & García-Díaz 2005; Mesa-Delgado et al. 2012) that satisfy existing observational constraints from high-resolution spectroscopy and *HST* imaging, and in which the abundances are adjusted so as to best reproduce the observed line spectrum. The second method has the great advantage that it can account for the fact that the physical conditions may vary greatly, even *within* the zone of emission of a single ion. It also obviates the need for the use of ICFs since the full ionization structure of each element is calculated from first principles, using the CLOUDY microphysics code (Ferland et al. 1998). On the other hand, it also has various disadvantages. The models are computationally costly to calculate, and there is no guarantee that a model will be found that can simultaneously reproduce all of the observations. In the case of several models that can each reproduce most, but not all, of the observed emission line spectrum, there is no obvious criterion for deciding which is 'best'.

³ The cusp is adopted as the area enclosed within $(X, Y) = (13-15, 2-4)$ while the background is two regions below $(5-6, 5-8)$ and $(18-19, 9-12)$

above left of the proplyd; in Argus coordinates with point of origin the bottom right corner of the map (cf. Fig. 3).

4.1 Empirical abundance determination

The density and temperature distributions derived in Section 3.3 were used in conjunction with the dereddened line ratio maps to compute ionic and elemental abundance maps relative to hydrogen. The IONIC task in IRAF v2.14 was implemented for this purpose. The [O III] $\lambda 4363/\lambda 4959$ temperature and [S II] $\lambda 4069/\lambda 6731$ density maps were adopted for the calculation. The $\lambda\lambda 6716+6731$, 3726+3729, 3967, 6584, 7136, 4959, 6312 forbidden line ratios relative to $H\beta$ were, respectively, used to compute S^+/H^+ , O^+/H^+ , Ne^{2+}/H^+ , N^+/H^+ , Ar^{2+}/H^+ , O^{2+}/H^+ and S^{2+}/H^+ ratios. Background-subtracted line intensities were adopted over the HST 10 area bounded by the 80 per cent $H\alpha$ isophote.

Mean values of the ionic and total abundances for the HST 10 cusp area and for the local M42 background from spaxels clear of HST 10 emission are included in Table 3, where they are compared with the values from our best-fitting photoionization model described in Section 4.2.

4.1.1 Oxygen and sulphur

The abundances for oxygen and sulphur should be well determined as they do not require the use of ICFs. The amounts of O^{3+} and S^{3+} in Orion are estimated to be infinitesimally small. The results are shown in Fig. 8. The O/H abundance map is the sum of the O^+/H^+ and O^{2+}/H^+ ratios. The abundance of oxygen is dominated by that of O^+ which makes up ≈ 80 per cent of O/H in the proplyd's cusp. On the other hand, S^{2+} makes up ≈ 80 per cent of S/H over the

same area. Similar percentages apply in the tail area where again the dominant oxygen ion is O^+ and the dominant sulphur ion is S^{2+} . The oxygen abundance is $12+\log(O/H) = 8.63 \pm 0.08$ and the abundance of sulphur is 6.75 ± 0.09 over HST 10. These are the mean values for the cusp area where the background subtraction is considered to be more accurate than over the fainter tail. For the local M42 background we find 8.56 ± 0.03 and 6.98 ± 0.10 for oxygen and sulphur, respectively.

The oxygen abundance derived for the local Orion nebula is practically the same as the one measured in the vicinity of LV 2 right within the Trapezium (Tsamis et al. 2011), and to the one measured by Esteban et al. (2004) at a position southwest of that (for a temperature fluctuation parameter, $t^2 = 0$). Sulphur in HST 10 is 0.2 dex lower than in its Orion vicinity or just 0.1 dex lower than the Orion nebula benchmark value based on IR data (Rubin et al. 2011); this indicates that sulphur is not depleted on to dust grains to a degree substantially greater than in the Orion nebula.

The gas-phase abundance of oxygen in the proplyd is not significantly different from that in the Orion H II region, and this is borne out from the model results too. Oxygen in HST 10 is equal within the uncertainties to the solar photospheric value of 8.69 ± 0.05 (Asplund et al. 2009), and to the mean value in B-type Orion OB1 stars of 8.74 ± 0.04 (Simón-Díaz 2010; Nieva & Simón-Díaz 2011). Adding the small amount of oxygen (~ 8.08) estimated to be in grains would bring the total oxygen at ~ 8.74 , which is equal to the protosolar value (Asplund et al. 2009) or the value based on B-stars in the solar vicinity (Nieva & Przybilla 2012).

4.1.2 Helium

The He^+/H^+ abundance ratio, which represents a measure of He/H in low ionization objects such as HST 10 where He^{2+} is non-existent, was measured using the $\lambda 4471$ and $\lambda 5876$ lines, adopting effective recombination coefficients from Smits (1996) and correcting for the effects of collisional excitation using the formulae in Benjamin, Skillman & Smits (1999). In both maps we measure a decreased helium abundance over the cusp of the proplyd by ~ 13 per cent compared to the Orion background (Fig. 8). Unrealistically low densities and/or high temperatures over HST 10 need to be assumed before the He abundance begins to approximate that in the background nebula; therefore the discrepancy is not due to the adopted physical conditions.

In our model of HST 10's cusp discussed in Section 4.2 this behaviour can be explained as it is shown that in HST 10 the ionization of He lags considerably behind that of hydrogen. At the ionization front, defined where H is 50 per cent ionized, He is only about 10 per cent ionized. The He abundance maps therefore show that significant amounts of neutral helium are present in the H^+ layer of the proplyd, thus rendering the direct abundance determination of He inaccurate. Using the ICF (equation 2 in Tsamis et al. 2011) by Peimbert, Torres-Peimbert & Ruiz (1992) to account for He^0 we estimate an upward correction of 14 per cent for He in the cusp (the correction for He^0 in the nebula is just 2 per cent). The resulting value agrees with that obtained from our detailed model according to which $He/H = 0.095$ in the cusp. This brings the HST 10 and Orion nebula helium abundances into very good agreement.

4.1.3 Nitrogen and neon

The abundances of nitrogen and neon were computed using the ICFs by Kingsburgh & Barlow (1994) and Peimbert & Costero (1969).

Table 3. Gas-phase ionic and element abundances for HST 10 and the Orion nebula (in decimal units for He and units of $\log H = 12$ for the other species).

Ion ratio	M42 $t^2 = 0$ (this work)	HST 10 $t^2 = 0$ (this work)		
C^{2+}/H^+ ^a	8.40 ± 0.04	8.06 ± 0.14		
N^+/H^+	7.00 ± 0.02	7.37 ± 0.03		
O^+/H^+	8.13 ± 0.05	8.45 ± 0.08		
O^{2+}/H^+	8.38 ± 0.03	7.98 ± 0.08		
O^{2+}/H^+ ^a	8.54 ± 0.05	–		
Ne^{2+}/H^+	7.58 ± 0.02	7.16 ± 0.05		
S^+/H^+	5.51 ± 0.08	5.90 ± 0.06		
S^{2+}/H^+	6.95 ± 0.13	6.68 ± 0.11		
Cl^{2+}/H^+	5.02 ± 0.07	4.78 ± 0.06		
Ar^{2+}/H^+	6.24 ± 0.05	6.18 ± 0.05		
Ar^{3+}/H^+	4.56 ± 0.18	–		
Fe^{2+}/H^+	5.72 ± 0.05	5.59 ± 0.05		
Element ratio	M42 $t^2 = 0$ (this work)	HST 10 $t^2 = 0$ (this work)	HST 10 model (this work)	M42 ^b $t^2 = 0.022$
He/H	0.097 ± 0.003	0.095 ± 0.012	0.095	0.097
C/H	8.59 ± 0.04	8.66 ± 0.14	8.48	8.42 ± 0.02
N/H	7.46 ± 0.04	7.48 ± 0.05	7.62	7.73 ± 0.09
O/H	8.56 ± 0.03	8.63 ± 0.08	8.63	8.65 ± 0.03
Ne/H	7.78 ± 0.04	7.75 ± 0.04	8.16	8.05 ± 0.07
S/H	6.98 ± 0.10	6.75 ± 0.09	6.75	7.22 ± 0.04
Cl/H	5.07 ± 0.08	4.90 ± 0.08	5.00	5.46 ± 0.04
Ar/H	6.27 ± 0.04	6.25 ± 0.05	6.30	6.62 ± 0.05
Fe/H	6.08 ± 0.05	5.70 ± 0.05	5.78	5.99 ± 0.10

^aValue derived from RLs.

^bFrom Esteban et al. (2004) for a nebula position 25 arcsec south and 10 arcsec west of θ^1 Ori C (≈ 45 arcsec north-west from HST 10).

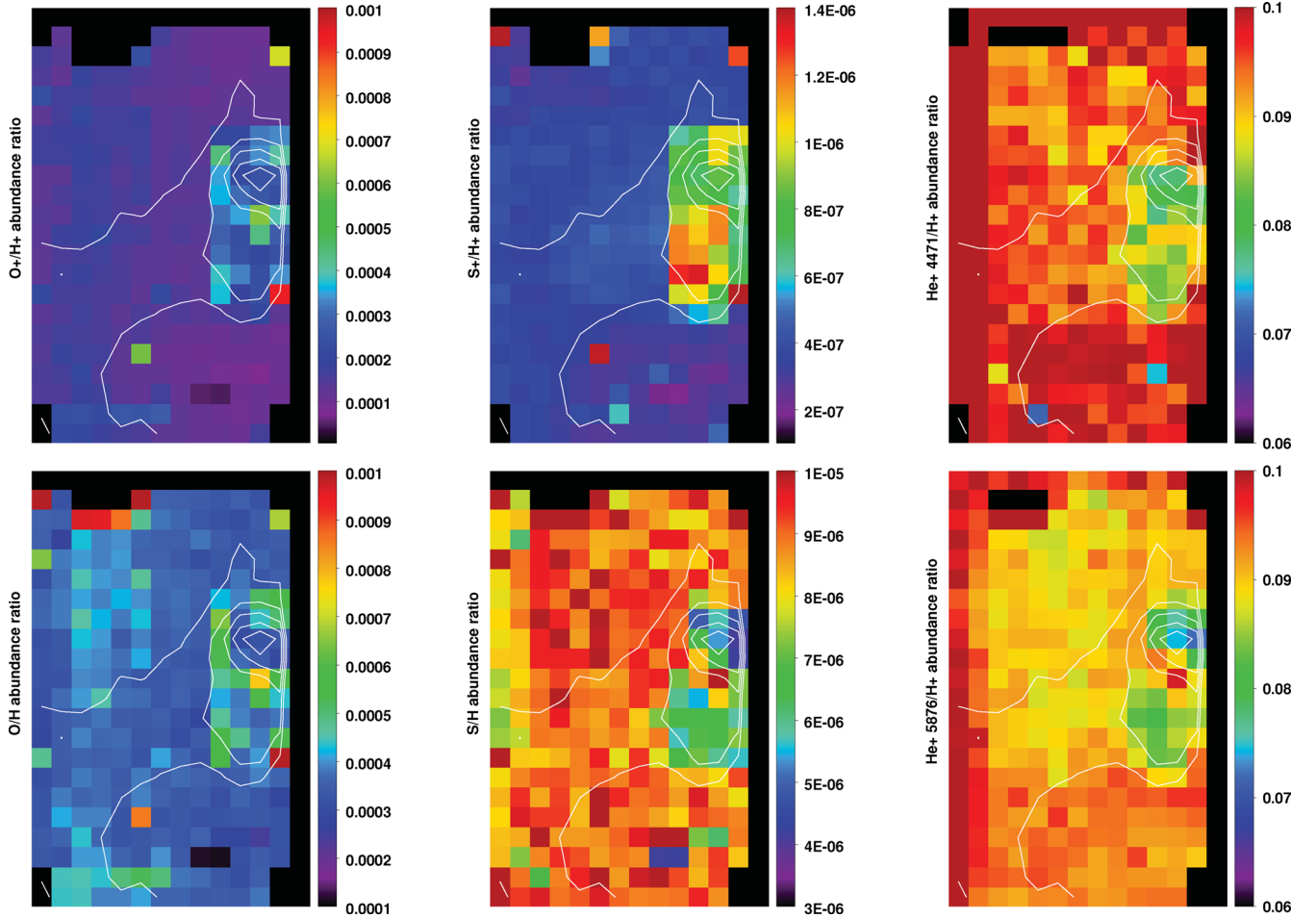


Figure 8. Abundance maps for O^+/H^+ , O/H , S^+/H^+ , S/H and He^+/H^+ (see the text for details).

For both elements we find no difference in their abundances between the proplyd and the nebula when comparing the ICF method determinations. However, our model of HST 10 (Section 4.2) calls for a nitrogen abundance 0.14 dex higher than the one derived using the classic ICF, where $N/O = N^+/O^+$. According to the model, the N abundance in HST 10 is 0.2 dex lower than in B-type Ori OB1 stars (Nieva & Simón-Díaz 2011).

There is a substantial difference between the Ne abundance for HST 10 derived by the ICF method and the one obtained by the simulations. Our best-fitting model (Section 4) has a neon abundance of $12 + \log(\text{Ne}/\text{H}) = 8.16$, i.e. 0.4 dex higher than the one derived using the classic ICF which calls for $\text{Ne}/\text{O} = \text{Ne}^{2+}/\text{O}^{2+}$. The ICF used to convert $\text{Ne}^{2+}/\text{H}^+$ to Ne/H for HST 10 is already fairly large ($= 3.9$) on account of the low ionization degree of the proplyd. A similar discrepancy was noted by Simón-Díaz & Stasinska (2011) whose model of the Orion nebula position observed by Esteban et al. (2004) calls for a neon abundance 0.2 dex higher than the one derived using the classic ICF. On the other hand, in models of the much higher ionization degree 30 Doradus region, Tsamis & Péquignot (2005) obtained neon abundances that were only ~ 20 per cent higher than those based on the ICF method.

These results indicate that the ICF method will systematically underestimate Ne/H in the low ionization conditions prevailing in structures such as HST 10. Our model supports a neon abundance for HST 10 slightly higher than the photospheric solar value of 7.93 ± 0.10 by Asplund et al. (2009). If the Orion nebula benchmark abun-

dance for neon of 8.01 ± 0.01 from the temperature and reddening insensitive IR $[\text{Ne II}]$ and $[\text{Ne III}]$ *Spitzer* measurements is adopted (Rubin et al. 2011), then HST 10 is slightly neon-rich compared to the Orion nebula (by 0.15 dex). However, compared to B-type stars in Orion (Nieva & Simón-Díaz 2011) the neon abundance from the model is equal within the uncertainties.

4.1.4 Carbon

The abundance of carbon rests on the determination based on the $\text{C II } \lambda 4267$ RL (Fig. 9) adopting the effective recombination coefficients of Davey et al. (2000), which include both radiative and dielectronic processes. The presence of C^+ was corrected for adopting $\text{ICF}(\text{C}) = \text{O}/\text{O}^{2+}$ (Kingsburgh & Barlow 1994), necessarily computed here from collisionally excited lines (since oxygen RLs are not detected from HST 10). The $\text{C II } \lambda 6578$ line, which was also detected, is dependent on optical depth effects and/or pumped by fluorescence and its diagnostic value is limited.

No difference is found in the abundance of carbon between the proplyd and the nebula; however the carbon abundance in our model of HST 10 is slightly lower than the one derived by the ICF method; this probably exposes the shortcomings of the latter when applied to HST 10. The carbon abundance from our preferred model is only ~ 0.1 dex higher than that in B-type stars in Orion (Nieva & Simón-Díaz 2011) or in a larger sample of B stars within 500 pc from the Sun (Nieva & Przybilla 2012).

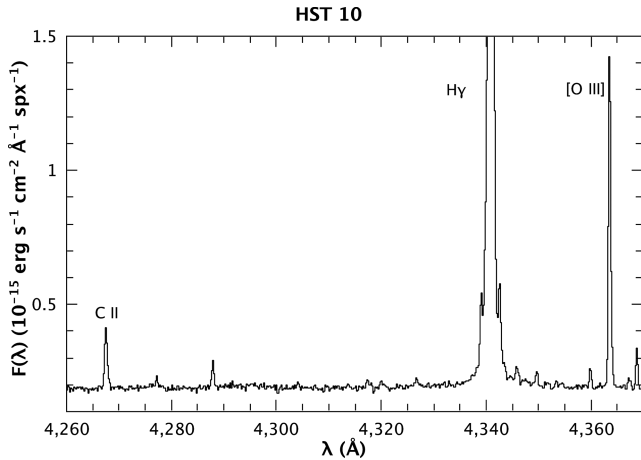


Figure 9. Spectrum of HST 10 extracted over the cusp area (not corrected for reddening), showing the C II $\lambda 4267$ RL and the [O III] $\lambda 4363$ collisional line.

4.1.5 Chlorine, argon and iron

The abundances of chlorine and argon were computed with the use of ICF equations (21) and (22) by Izotov et al. (2006), and are respectively almost equal in the proplyd and the nebula. The abundances of these species obtained by our detailed model of HST 10 do not differ substantially from those derived by the ICF method.

The abundance of iron was computed from the [Fe III] $\lambda 5270$ line (Fig. 4) using a 25-level atomic model for Fe III (cf. Tsamis & Walsh 2011). The resulting $\text{Fe}^{2+}/\text{H}^+$ ratio was converted to Fe/H using the ICF by Rodríguez & Rubin (2005). The proplyd shows an iron abundance ~ 0.3 dex lower than the Orion nebula. The abundance of gas-phase Fe therefore appears to be significantly different between the proplyd and the nebula. The abundance of Fe in the Orion H II region is already depleted relative to solar by ~ 1.5 dex (Esteban et al. 2004; Tsamis et al. 2011) due to iron being locked in dust grains. As a result, more iron appears to be trapped in condensates in HST 10 than is the case for the local Orion nebula. The observed overall depletion in HST 10 is, however, not as strong as the one encountered in LV 2 (-2.5 dex relative to solar for LV 2's rest velocity component); that proplyd shows an Fe depletion similar to those observed in the cold interstellar medium (Tsamis et al. 2011; Tsamis & Walsh 2011).

4.1.6 The recombination/collisional-line abundance anomaly

We note that the long-standing RL/CL abundance anomaly (e.g. Tsamis et al. 2003; Esteban et al. 2004; García-Rojas & Esteban 2007; Mesa-Delgado & Esteban 2010; Simón-Díaz & Stasińska 2011; Tsamis et al. 2011; Mesa-Delgado et al. 2012) cannot be quantified in the case of HST 10. This is defined as the difference in ionic abundances derived from metallic RLs compared to those derived from collisionally excited lines emitted from the same ions (the so-called abundance discrepancy factor, ADF). In the Orion nebula the ADF for O^{2+} is ~ 0.1 – 0.3 dex (e.g. Mesa-Delgado et al. 2012, and references therein). On the other hand, the ADF is practically zero in the cusps of the LV 2 (Tsamis et al. 2011) and 177–341 proplyds (Mesa-Delgado et al. 2012). The accurate removal of the contaminating Orion nebula emission and the adoption of realistic temperatures and densities for the proplyd doubly ionized carbon and oxygen regions are the most likely causes of these encouraging results.

In the case of HST 10, with respect to carbon only C II $\lambda 4267$ is observed while its CL counterpart C III] $\lambda 1908$ is not available. RLs of O^{2+} , which are weaker than C II $\lambda 4267$, are not detected in the background-subtracted spectra of HST 10 tallying with the overall low ionization degree of the proplyd. The intensity of the strongest predicted O II $\lambda 4649.13$ V1 multiplet line is dominated by the Orion nebula emission.

We derived the mean O^{2+}/H^+ ratio for ~ 20 spaxels of the background nebula from the summed intensities ($I = 0.3265$ in units where $\text{H}\beta = 100$) of five O II V1 RLs (4638.86, 4649.13, 4650.84, 4661.63, 4676.24 Å), adopting effective recombination coefficients from Storey (1994). Comparing with the abundance derived from the [O III] $\lambda 4959$ CL we obtain an $\text{ADF}(\text{O}^{2+}) \simeq 0.16$ dex for the Orion nebula within our field of view (Table 3).

4.2 Photoevaporation models of HST 10

We have calculated dynamic photoevaporation models of HST 10 with the photoionization code CLOUDY v10.0 (Ferland et al. 1998), using the procedure outlined in Mesa-Delgado et al. (2012). The parameters for the models are shown in Table 4. As compared with proplyd 177–341 (HST 1), which was modelled in Mesa-Delgado et al. (2012), HST 10 receives a roughly 10 times smaller ionizing flux, F , and is roughly twice as large. Since $F \propto n^2 r$ for recombination-dominated photoevaporation flows (Henney 2001), this implies that the densities in HST 10 should be $\simeq 5$ times smaller and the ionization parameter $\simeq 2$ times smaller than in HST 1. Fig. 10 shows the variation of physical variables along a representative radial cut through our best-fitting model described below. The ionizing flux was adjusted until the model reproduced the correct H β flux for the proplyd cusp. Assuming an ionizing flux of $9.0 \times 10^{48} \text{ s}^{-1}$ for θ^1 Ori C and no intervening absorption, this implies a true distance from the Trapezium of about twice the projected distance, or an inclination of about 60° from the plane of the sky. This is consistent with the appearance in the *HST* images and with the kinematics measured from Keck spectra.

Two different abundance sets were used in the models. The first (Model A) is the Orion gas-phase abundance set, as determined by Esteban et al. (2004) for a temperature fluctuation parameter, $t^2 > 0$. The second (Model B) is a set of abundances starting from those determined in this paper using the semi-empirical method and assuming that $t^2 = 0$ (Table 3). The emissivity structure of Model B shown in the middle panel of Fig. 10 can be compared

Table 4. Input parameters for example physical models of 182–413.

Stellar spectrum:	$T_* = 39\,000 \text{ K}$
Simón-Díaz et al. (2006)	$\log g = 4.1$
	$L_* = 2.04 \times 10^5 L_\odot$
Projected distance to θ^1 C:	57 arcsec (O'Dell 1998)
Inclination angle:	30 deg (Henney & O'Dell 1999)
Derived physical distance to θ^1 C:	$7.51 \times 10^{17} \text{ cm}$
Ionizing flux at proplyd:	$\Phi_{\text{H}} = 1.27 \times 10^{12} \text{ cm}^{-2} \text{ s}^{-1}$
Ionization front radius:	$r_0 = 3.7 \times 10^{15} \text{ cm}$ (Henney & O'Dell 1999)
Gas-phase abundances:	Model A: Esteban et al. (2004)
($12 + \log_{10} z/\text{H}$)	He 10.99, C 8.42, N 7.73, O 8.65
	Ne 8.05, S 7.22, Cl 5.46, Ar 6.62, Fe 6.0
	Model B: ad hoc for HST 10
	He 10.98, C 8.48, N 7.62, O 8.63
	Ne 8.16, S 6.75, Cl 5.0, Ar 6.30, Fe 5.78
Dust composition:	Standard Orion (Baldwin et al. 1991)

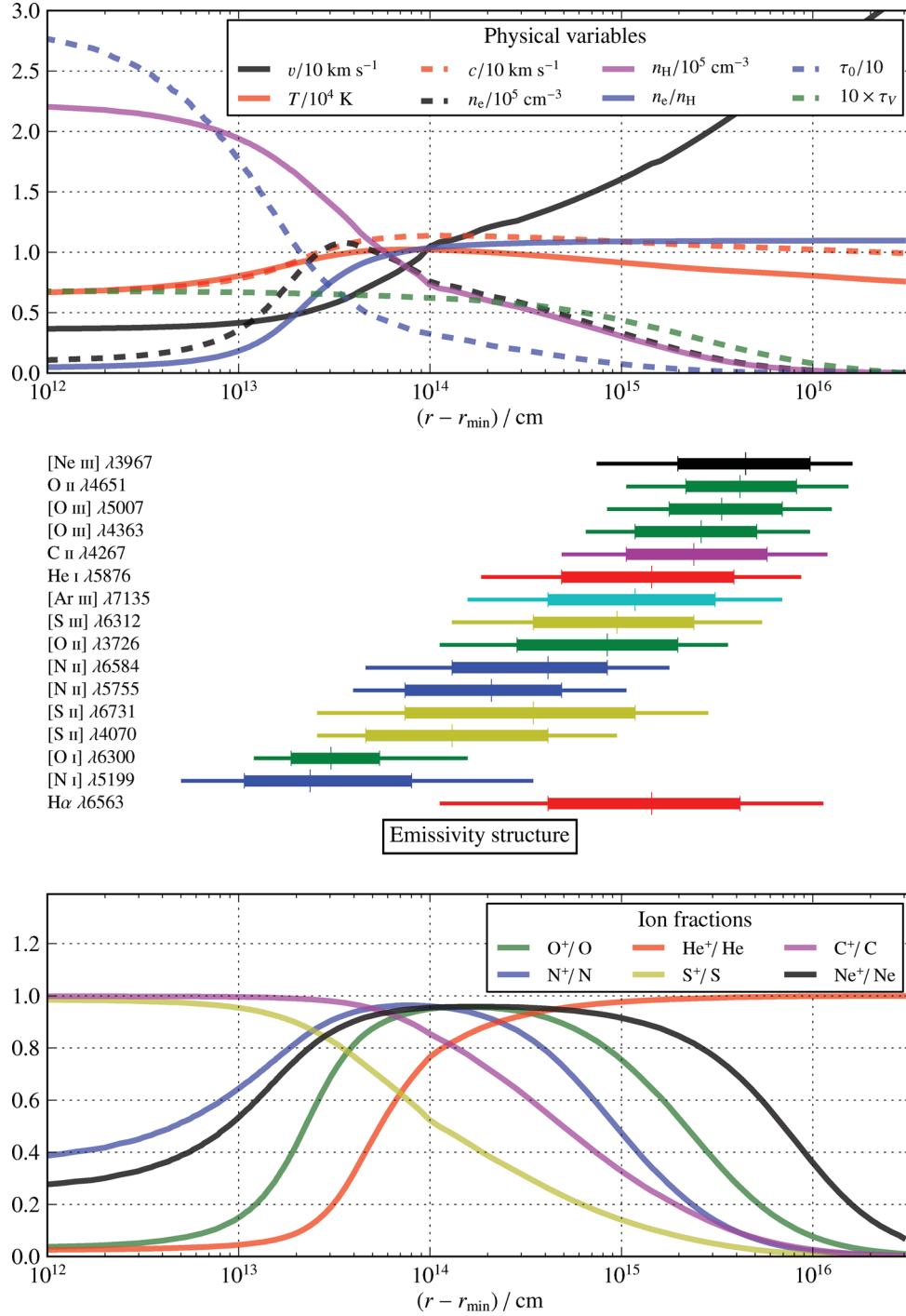


Figure 10. Model structure as a function of radius along a line at an angle of $\theta = 30^\circ$ to the symmetry axis for our photoevaporation Model B. The radius r is measured from the centre of curvature of the ionization front, shown on a logarithmic scale with respect to the deepest point in the model, $r_{\min} = 3.5 \times 10^{15} \text{ cm}$. The ionization front is at $r - r_{\min} \simeq 2 \times 10^{13} \text{ cm}$. Upper panel: physical variables. Gas velocity, v , in units of 10 km s^{-1} (solid black line). Gas temperature, T , in units of 10^4 K (solid red line). Isothermal sound speed, c , in units of 10 km s^{-1} (dashed red line). Electron density, n_e , in units of $7 \times 10^4 \text{ cm}^{-3}$ (dashed black line). Total hydrogen density, n_H , in units of $7 \times 10^4 \text{ cm}^{-3}$ (solid purple line). Electron fraction, n_e/n_H (solid blue line). Lyman limit optical depth, τ_0 (divided by 10), measured from the ionizing source, which is located off the graph at $r = 6.32 \times 10^{17} \text{ cm}$ (dashed blue line). Visual dust extinction optical depth, τ_V (times 10) (dashed green line). Middle panel: emissivity structure for selected emission lines. For each emission line, the vertical line indicates the median emission radius (for which half the total line flux is emitted from gas at smaller radii and half from gas at larger radii), while the thick horizontal bar shows the range between the 25th and 75th quartiles of the emission, and the thin horizontal bar shows the range between the 10th and 90th quartiles of the emission. Lower panel: ionization structure. The fraction of singly ionized ions for a selection of elements is shown by differently coloured lines, as indicated in the key. Results for other angles and abundance sets are similar, except with densities that scale roughly as $\cos^{1/2}\theta$, and temperatures that tend to decrease slightly with θ and with increased metal abundances.

with the emission line maps presented in Section 3.1. The model predicts that the lower ionization lines, such as those of [O I], [S II] and [N II], should show emission that is offset towards the centre of the proplyd with respect to the H α emission by $1\text{--}2 \times 10^{15}$ cm ($0.5\text{--}1$ spaxel), which is very similar to what is observed (Fig. 3). In contrast, the highest ionization lines, such as those of [O III] and [Ne III], should be offset away from the centre of the proplyd by a similar amount, which is again consistent with the observed spatial distributions (Fig. 4). The model also makes more detailed predictions about the emission distributions, such as the fact that for both [O III] and [N II] the nebular collisional lines should be weighted towards slightly larger radii than the auroral lines, a consequence of the temperature and density gradients within a given ionization zone. There is some evidence for this from our emission maps, but the predicted magnitude of the effect ($0.1\text{--}0.5$ spaxel) is at the limit of what can be discerned with our spatial resolution.

The spectrum from the two models is shown in Fig. 11. It can be seen that the Esteban et al. (2004) abundances produce large

discrepancies between the observed and predicted line fluxes (panel ‘a’). In particular, all sulphur, argon and chlorine lines are too strong in the model by a factor of 3 to 10, whereas [O III] $\lambda 4363$ is too weak, indicating that the model temperature is too low in the highly ionized regions. On the other hand, the [N II] lines are well reproduced by this model.

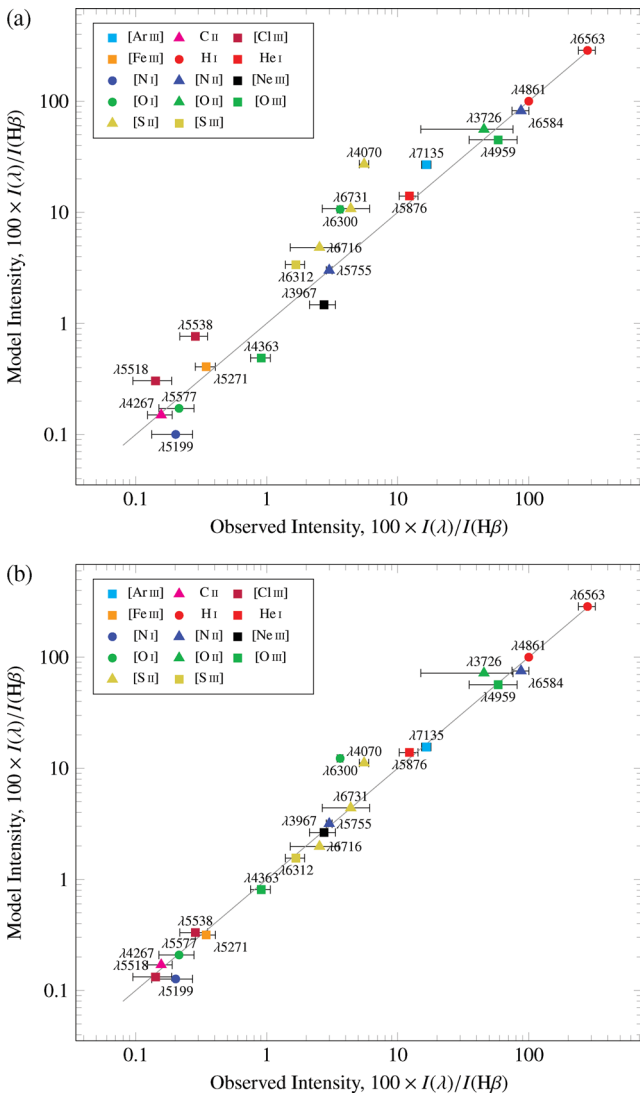
The situation is much improved by using the Model B abundances (Fig. 11b), although some discrepancies remain. The sulphur line fluxes are now in good agreement with the observations, with the notable exception of the [S II] $\lambda 4069$ auroral line, which is still three times too strong in the model. The [O I] $\lambda 6300$ shows a similar behaviour, being too strong by a factor of 4. These two lines, together with the [N II] $\lambda 5755$ auroral line, show the strongest brightness contrast between the proplyd and the background nebula, and hence are measured with a relatively small uncertainty, making the disagreement highly significant. The [N II] lines do not agree with Model B so well as they do with Model A, but the agreement is still fair given the uncertainties. The remaining disagreement is with the [N I] $\lambda 5199$ line, but this is to be expected since the line arises through fluorescence in neutral gas (Ferland et al. 2012), whereas the model extends only to a hydrogen ionization fraction of 0.1 and therefore misses part of the zone emitting [N I].

The two most discrepant lines, [O I] $\lambda 6300$ and [S II] $\lambda 4069$, have a considerable overlap in their zones of emission (Fig. 10). All the [O I] emission, and roughly half of the [S II] $\lambda 4069$, arises from partially ionized gas at the ionization front itself. If the temperature of this gas is overestimated in the models, then the discrepancy could be explained. With the standard Orion dust properties that we are using in all the models (Baldwin et al. 1991), photoelectric emission from dust grains provides up to 15 per cent of the heating in this zone. Since there is some evidence that dust is depleted in the proplyd flows (García-Arredondo, Henney & Arthur 2001), it may be feasible to reduce this heating, which is an avenue that will be explored in follow-up work.

Notwithstanding these discrepancies, we consider that Model B for HST 10 has a much better defined O/H abundance than our model for HST 1. This is because the [O II] $\lambda 3726$, $\lambda 3729$ lines were not measured from that proplyd due to their extreme weakness from collisional quenching in the prevailing high-density conditions. With regard to chlorine and argon our preferred model for HST 10 does not support the high ($r^2 > 0$) abundances for these elements in Orion obtained by Esteban et al. (2004). As a noble gas, Ar is not expected to be depleted on grains (even though it can form clathrates in ices; see e.g. Mousis et al. 2010), and therefore the 0.3 dex difference with the Orion value by Esteban et al. is perhaps significant.

4.3 Comparison with other proplyds

HST 10 is the third proplyd to be subject to a detailed abundance analysis, following earlier studies of LV 2 (Tsamis et al. 2011; Tsamis & Walsh 2011) and HST 1 (Mesa-Delgado et al. 2012). These proplyds cover a broad range in size and in separation from the Trapezium stars (see Fig. 12), and derived physical parameters for the three proplyds are summarized in Table 4. HST 10 shows a significantly lower ionization parameter than the two closer-in proplyds, which is reflected in its emission line spectrum that is relatively stronger in low ionization lines. Despite these differences, the estimated mass-loss rate from the ionized cusp is very similar for all three proplyds, being of the order of $2 \times 10^{-7} M_{\odot} \text{ yr}^{-1}$. These values are somewhat lower than earlier estimates (e.g. Henney &



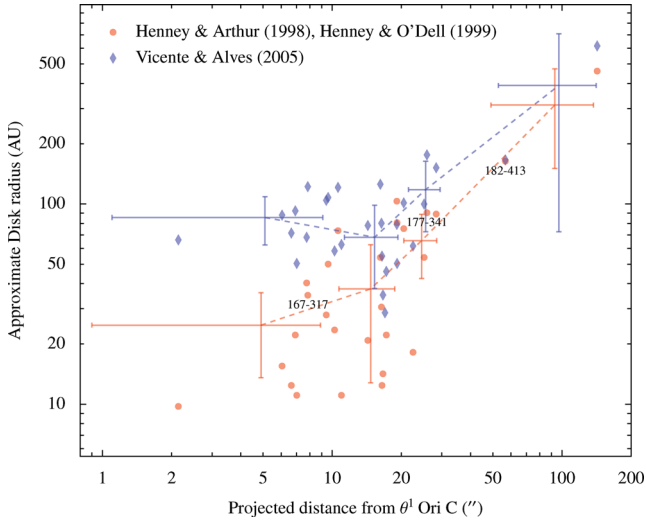


Figure 12. Variation with projected distance from the ionizing star of circumstellar disc sizes in the Orion proplyds. Blue symbols show empirical determinations from Vicente & Alves (2005); red symbols show results from fitting photoevaporation models (Henney & Arthur 1998; Henney & O'Dell 1999). The three proplyds that have been subject to gas-phase abundance studies are marked: 167–317 (LV 2), 177–341 (HST 1) and 182–413 (HST 10). In only a few cases, such as HST 10, is the molecular disc size directly measured, for the rest it is assumed to be half the size of the ionization front. Note that the Vicente & Alves (2005) methodology significantly overestimates the true size of the ionization front (and by extension the enclosed disc) for small proplyds or those that are much brighter than the surrounding nebula, which tend to be those found close to the ionizing star. Error bars show the mean and standard deviation measured in four broad spatial bins.

O'Dell 1999; Henney et al. 2002), which is partly because we are neglecting the contribution of mass loss through the proplyd tail.

Table 5 also shows estimates for the mass and radius of the embedded circumstellar accretion disc, which is the reservoir of mass in the proplyds. All three proplyds show very similar sub-mm fluxes (Mann & Williams 2010), of the order of 20 mJy once the contribution from ionized free-free emission has been subtracted. However, conversion of this flux to a gas mass requires knowledge of the dust opacity per unit gas mass and dust temperature, both of which have large uncertainties (Williams & Cieza 2011). The values in the table are calculated on the assumption that the dust temperature is the effective temperature of a disc in radiative equilibrium with the bolometric flux from the Trapezium stars,⁴ and using the opacity recommended by Beckwith et al. (1990) of $0.1(\nu/1000 \text{ GHz}) \text{ cm}^2 \text{ g}^{-1}$, giving $0.034 \text{ cm}^2 \text{ g}^{-1}$ at $880 \mu\text{m}$. It must be emphasized that the derived masses are highly uncertain, since the opacity could be up to 20 times smaller if substantial grain growth up to cm-sized bodies has occurred (D'Alessio, Calvet & Hartmann 2001). Evidence for grain evolution has been found in the case of silhouette discs projected on to the Orion nebula (Miotello et al. 2012). The masses given in the table are at least six times smaller than the ‘minimum mass solar nebula’ (Weidenschilling

⁴ Although the sub-mm emission is optically thin, the discs are likely to be optically thick at mid-infrared wavelengths where they emit the bulk of their radiation. We assume that 50 per cent of the bolometric radiation from the Trapezium stars is absorbed in the ionized cusp of the proplyd, and that 50 per cent of the remainder is absorbed in the neutral photoevaporation flow, so that only 25 per cent reaches the surface of the disc.

1977), which is the mass within 30 au required to account for the composition of the planets in the Solar system.

The nominal photoevaporation ages, $t_{\text{evap}} = M_d/\dot{M}$, are uncomfortably short compared with the estimated age of the photoionized nebula ($\geq 10^5$ yr; see discussion in section 8.2.2 of Henney & O'Dell 1999), but would come into agreement if the masses were increased by a factor of 5–10, which cannot be ruled out (see the previous paragraph). Interestingly, the evaporation age t_{evap} increases with increasing distance from the Trapezium. Given that t_{evap} should be roughly equal to the elapsed time since the disc photoevaporation commenced (Johnstone et al. 1998), this is the opposite of what would be expected from a naive model of a roughly spherical H II region, in which case proplyds at greater distances from the ionizing star would have entered the H II region more recently. However, evidence from both observations (O'Dell et al. 2009) and numerical simulations (Arthur et al. 2011) point to the continued survival of dense clumps of molecular gas well inside the apparent boundary of the H II region, in which case it is reasonable that the proplyds closest to the Trapezium might have been shielded from ultraviolet radiation until relatively recently.

The abundance results for the three proplyds are rather disparate, so that it is very hard to see a consistent trend in the results. For instance, although we find a roughly solar oxygen abundance for HST 10 from both empirical analysis and photoevaporation model fitting of [O II] and [O III] CLs, the oxygen abundance was found to be $\simeq 2$ times solar in LV 2 (Tsamis et al. 2011) from a purely empirical analysis, with no discrepancy between CLs and RLs and with both O^+ and O^{2+} stages observed (in [O II] CLs for the former and [O III] CLs and O II RLs for the latter). In contrast, Mesa-Delgado et al. (2012) found oxygen to be about $\simeq 0.4$ times solar in HST 1, from photoevaporation model fitting of O^{2+} ([O III] CLs), whereas ‘the empirical’ analysis of O II RLs is consistent with the [O III] CLs. Given the wide range of characteristic ionization and density found in the three proplyds (Table 5), it is possible that systematic errors in our abundance analysis might be contributing to this wide spread. In order to rule out any such effects, it would be very worthwhile to carry out a similar analysis on a sample of proplyds that all have similar ionization parameters and densities.

5 SUMMARY AND CONCLUSIONS

We have presented integral field spectroscopic observations of the HST 10 proplyd found in the Orion nebula obtained with the VLT and produced a collection of sub-arcsec resolution monochromatic images of the proplyd and the surrounding nebula. The IFS observations facilitated the isolation of the proplyd's intrinsic emission from the emission of the Orion nebula. As a result, the reddening, ionization degree, electron temperature and density, as well as the ionic and elemental abundances for several species, have been derived for the proplyd and for the local Orion neighbourhood.

The abundances of oxygen and sulphur have been derived for HST 10 without resort to ICFs and are not significantly different from those measured in the local Orion nebula. This conclusion is also borne out from a detailed photoevaporation model of the proplyd. The abundances of carbon, oxygen and neon in HST 10 are practically the same as those in B-type stars in Orion.

Based on our model, the abundance of neon in HST 10 is 0.2 dex higher than the photospheric solar value by Asplund et al. (2009) and 0.15 dex higher than the IR *Spitzer* value by Rubin et al. (2011). Compared to the model, the standard ICF method for

Table 5. Comparison of physical properties between HST 10 and two other well-studied proplyds.

	Units	Note	LV 2	HST 1	HST 10
Coordinate-based designation		1	167–317	177–341	182–413
<i>Relation to ionizing source</i>					
Projected distance, D'	arcsec	2	7.83	25.84	56.7
Inclination, i	°	3	50	70	150
True distance, D	pc	4	0.022	0.059	0.242
<i>Ionized cusp</i>					
Ionization front radius, r_0	au	5	53.	136.	247.
Peak electron density, n_0	10^6 cm^{-3}	6	2.0	0.4	0.1
Ionization parameter		7	0.012	0.008	0.002
Cusp mass-loss rate, \dot{M}	$10^{-7} M_{\odot} \text{ yr}^{-1}$	8	2.6	2.5	2.1
<i>Molecular disc</i>					
Disc effective temperature, T_d	K	9	95	58	29
Disc mass, M_d	$10^{-3} M_{\odot}$	10	1.6	2.7	5.4
Disc radius, R_d	au	11	34	89	160
Evaporation age, t_{evap}	10^4 yr	12	0.6	1.1	2.6

Notes: (1) O'Dell & Wenn (1994). (2) Angular separation from θ^1 Ori C (O'Dell 1998). (3) Inclination of proplyd axis to line of sight estimated from kinematic studies of the velocity–ionization correlation in emission lines from the cusp (Henney & O'Dell 1999; Henney et al. 2002). Proplyds with $i > 90^\circ$ have their head pointing away from the observer. (4) $D = D'/\sin i$. (5) Estimated from fitting evaporation models to the H α profiles of the cusps (Henney & Arthur 1998). (6) LV 2 from C III] density (Henney et al. 2002); HST 1 and HST 10 from model fitting (this paper and Mesa-Delgado et al. 2012). (7) $F/(n_0 c)$. (8) Calculated by integrating model mass fluxes over the area of the cusp. (9) Radiative equilibrium temperature, assuming that 25 per cent of the bolometric flux from θ^1 Ori C reaches the surface of the disc (see also Robberto, Beckwith & Panagia 2002). (10) Estimated from observed fluxes at 880 μm (Mann & Williams 2010) after subtracting the contribution from ionized free–free emission, assuming optically thin dust emission with opacity $\kappa_\nu = 0.034 \text{ cm}^2 \text{ g}^{-1}$ and dust temperature equal to the effective temperatures derived above. (11) Directly estimated from *HST* images for HST 10. For LV 2 and HST 1, we assume $r_d = 0.65r_0$ (see Fig. 12). (12) Nominal mass-loss time-scale: M_d/\dot{M} .

neon significantly underestimates the element's abundance in HST 10 by 0.4 dex. This indicates that the classic ICF method applied to optical [Ne III] emission lines can systematically underestimate Ne/H in low ionization Orion structures.

Iron in HST 10 is depleted by ~ 0.3 dex compared to the Orion nebula as a result of a higher percentage of Fe trapped in solids than in the ambient H II region.

Based on our derived mass-loss rates and observed disc fluxes at mm wavelengths, the photoevaporation ages of the HST 10, LV 2 and HST 1 proplyds are very short ($\simeq 10^4 \text{ yr}$) if one assumes 'standard' dust opacities (Beckwith et al. 1990), which is in conflict with lower limits on the age of the Orion nebula ($> 10^5 \text{ yr}$; Henney & O'Dell 1999). This implies that the mm dust opacity must be lower, possibly due to grain growth to cm sizes. The photoevaporation ages are found to be lowest for the proplyds closest to θ^1 Ori C, implying that these must have been shielded inside dense molecular gas until relatively recently, which is consistent with a highly irregular shape for the Orion nebula's ionization front, as has been suggested by recent observational (O'Dell et al. 2009) and theoretical (Arthur et al. 2011) studies.

In conclusion, our study of HST 10 – and those of LV 2 and HST 1 based on IFS techniques – shows that spatially resolved spectrophotometry accompanied by detailed simulations can be successfully applied to evaporating protoplanetary discs, helping to establish their chemical composition in Orion for the first time. These studies of the youngest component of this landmark region's ecosystem complement superbly the abundance analyses of ionized gas and stars in the same area.

ACKNOWLEDGEMENTS

We thank the VLT/FLAMES astronomers at ESO for supporting the service mode observations (programme 078.C-0247(A); PI: Tsamis). We also thank the referee Dr. C. R. O'Dell for a constructive report. This research made use of the NASA ADS digital library portal operated by the Smithsonian Astrophysical Observatory (SAO); ESO-MIDAS on Mac/OSX provided by the European Southern Observatory; IRAF, distributed by the National Optical Astronomy Observatory, which is operated by the Association of Universities for Research in Astronomy; SAOImage DS9 developed by SAO.

WJH and NFF acknowledge financial support from DGAPA-UNAM through project PAPIIT IN102012 and from a postdoctoral fellowship to NFF. AMD acknowledges support from Comité Mixto ESO-Chile and a Basal-CATA (PFB-06/2007) grant.

YGT acknowledges support from a Marie Curie intra-European Fellowship within the 7th European Community Framework Programme (project PIEF-GA-2009-236486), and thanks his collaborators at the Centro de Radioastronomía y Astrofísica in Morelia for their warm hospitality during his visit.

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